

Improving pulsar timing precision through superior time-of-arrival creation

J. Wang^{1,2,3}, J. P. W. Verbiest⁴, G. M. Shaifullah^{5,6}, I. Cognard^{7,8}, L. Guillemot^{7,8}, G. H. Janssen^{9,10}, M. B. Mickaliger¹¹, A. Possenti¹², and G. Theureau^{7,8,13}

¹ Ruhr-Universität Bochum, Fakultät für Physik und Astronomie, Astronomisches Institut (AIRUB), 44801 Bochum, Germany
e-mail: jun.wang.ucas@gmail.com

² Advanced Institute of Natural Sciences, Beijing Normal University, Zhuhai 519087, PR China

³ Fakultät für Physik, Universität Bielefeld, Postfach 100131, 33501 Bielefeld, Germany

⁴ Florida Space Institute, University of Central Florida, 12354 Research Parkway, Orlando, FL 32826, USA

⁵ Dipartimento di Fisica ‘G. Occhialini’, Università degli Studi di Milano-Bicocca, Piazza della Scienza 3, 20126 Milano, Italy

⁶ INFN, Sezione di Milano-Bicocca, Piazza della Scienza 3, 20126 Milano, Italy

⁷ Laboratoire de Physique et Chimie de l’Environnement et de l’Espace LPC2E CNRS-Université d’Orléans, 45071 Orléans, France

⁸ Station de radioastronomie de Nançay, Observatoire de Paris, PSL Research University, CNRS/INSU, 18330 Nançay, France

⁹ ASTRON, the Netherlands Institute for Radio Astronomy, Oude Hoogeveensedijk 4, 7991 PD Dwingeloo, The Netherlands

¹⁰ Department of Astrophysics/IMAPP, Radboud University Nijmegen, PO Box 9010, 6500 GL Nijmegen, The Netherlands

¹¹ Jodrell Bank Centre for Astrophysics, University of Manchester, Manchester M13 9PL, UK

¹² INAF – Osservatorio Astronomico di Cagliari, via della Scienza 5, 09047 Selargius (CA), Italy

¹³ Laboratoire Univers et Théories LUTH, Observatoire de Paris, PSL Research University, CNRS/INSU, Université Paris Diderot, 5 place Jules Janssen, 92190 Meudon, France

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ABSTRACT

Context. The measurement of pulsar pulse times-of-arrival (ToAs) is a crucial step in detecting low-frequency gravitational waves. To determine ToAs, we can use template-matching to compare each observed pulse profile with a standard template. However, using different combinations of templates and template-matching methods (TMMs) without careful consideration can lead to inconsistent results.

Aims. In pulsar timing array (PTA) experiments, distinct ToAs from the same observations can be obtained due to the use of diverse templates and TMMs. In other words, employing diverse approaches can yield different timing results and would thus have a significant impact on subsequent gravitational wave searches. In this paper we examine several commonly used combinations to analyze their effect on pulse ToAs.

Methods. We evaluated the potential impact of template and TMM selection on 13 typical millisecond pulsars within the European PTA. We employed pulsar timing methods, specifically the root mean square and reduced chi-square (χ_r^2) of the residuals of the best timing solution, to assess the outcomes. Additionally, we evaluated the system-limited noise floor (SLNF) for each pulsar at various telescopes operating around 1.4 GHz using frequency-resolved templates.

Results. Our findings suggest that utilizing data-derived and smoothed templates in conjunction with the Fourier-domain with Markov chain Monte Carlo TMM is generally the most effective approach, though there may be exceptions that require further attention. Furthermore, we determine that pulse phase jitter noise does not significantly limit the current precision of the European PTA’s timing, as jitter levels derived from other studies are much lower than the SLNF.

Key words. methods: data analysis – pulsars: general

1. Introduction

Since their discovery, millisecond pulsars (MSPs; [Backer et al. 1982](#)) have been known for their short spin periods and highly stable rotation characteristics. These unique properties have allowed for a diverse range of applications in testing fundamental physics, including high-precision tests of general relativity ([Zhu et al. 2015](#); [Kramer et al. 2021](#)), the exploration of the ionized interstellar medium ([Turner et al. 2021](#); [Liu et al. 2022](#)) and solar wind ([Tiburzi et al. 2019, 2021](#)), the development of pulsar-based timescales ([Hobbs et al. 2012, 2020](#)), constraints on the equations-of-state for neutron stars ([Lattimer & Prakash 2016](#); [Özel & Freire 2016](#)), and measurements of

the model-independent properties of the pulsar itself ([Verbiest et al. 2008](#)).

One of the most significant applications of MSP observations, however, lies in the detection and characterization of low-frequency (nanohertz) gravitational waves (GWs; see, e.g., [Verbiest et al. 2016](#); [Perera et al. 2019](#)). These waves represent fluctuations in the fabric of space-time caused by the propagation of GWs, as predicted by the theory of general relativity (see, e.g., [Misner et al. 1973](#)). Pulsar timing arrays (PTAs), which consist of an array of MSPs, are particularly sensitive to low-frequency GW signals. In this frequency range, the potential sources of GWs are likely an ensemble of supermassive black hole binaries, whose combined, stochastic signal gives rise to the

GW background (Burke-Spolaor et al. 2019). Typically, several factors can influence the precision of pulsar timing, and each source of noise has a unique impact (Verbiest & Shaifullah 2018). However, GWs leave a distinct quadrupolar signature in the pulsar timing residuals. A method based on the correlation between pairs of pulsars in the sky as a function of their angular separation, known as the Hellings and Downs curve (Hellings & Downs 1983), can be employed to identify this signature of GWs.

Pulsar timing analysis involves determining the differences, referred to as timing residuals, between the observed times-of-arrival (ToAs) and the predicted ToAs. These predictions are based on models that take into account, for example, the influence of the astrometry and ephemeris of the MSPs in the array, the Earth, and Solar System bodies. A precise and comprehensive model would result in accurate timing residuals. However, the precision of the timing, quantified by the root mean square (RMS) value of the timing residuals, is constrained by various sources of noise originating from the pulsar itself, the propagation path, or the observing system. To enhance timing precision and increase the likelihood of detecting GWs, researchers have explored and implemented various analysis methods and hardware upgrades in PTA observations and data analysis. Nevertheless, the impact of different data analysis pipelines at different telescopes has not been thoroughly investigated. To address this issue, Wang et al. (2022) propose a method for assessing the influence of different template options and template-matching methods (TMMs) using three pulsars. Here we briefly review some of its basic concepts. The Fourier phase gradient (PGS) determines the shift between two similar profiles by fitting for the gradient in the Fourier domain. The Fourier-domain with Markov-chain Monte Carlo (FDM) method is identical to the PGS in how it determines ToA values but applies a one-dimensional Markov chain Monte Carlo method to measure the ToA variance. The Gaussian interpolation shift (GIS) method calculates the discrete cross-correlation function between the profile and the template, and a Gaussian curve is then fitted to the resulting points to enable interpolation between each phase bin. For the construction of the single template, the observation with the highest signal-to-noise ratio (S/N) is selected to serve as the template and is excluded from the set of observations to be timed. For the added template, all observations are added up and aligned using the timing ephemeris of the data. For the smoothed template, the added profile is then refined using an undecimated Daubechies wavelet smoothing filter to produce the final smoothed template. The analytic template is generated by fitting von Mises functions to the added profile. A more detailed description of the TMMs and templates is provided in Wang et al. (2022).

Wang et al. (2022) conclude that data-derived and smoothed templates are typically preferable, and that the FDM method generally outperforms other methods. However, given the limited number of pulsars in Wang et al. (2022), these findings should be examined and verified with a larger sample of pulsars. More recently, Wang et al. (2023) delved into TMMs for pulsars, scrutinizing the shape of the pulse profile and exploring the impact of the S/N of both the template and the observation on the ToA uncertainty. They achieved this by examining pulsar pulse profiles obtained from the International Pulsar Timing Array (IPTA). Through an extensive analysis involving numerous simulation models and real data, Wang et al. (2023) discovered that while the ToA uncertainty computed with the FDM method is generally more reliable, there are instances where alternative methods yield better-behaved ToA uncertainties compared

to FDM. This occurs particularly in cases where the duty cycle of the pulse profile is long, the S/N of the observations is low, or the pulse profile lacks sharp features. However, since the study only focused on ToA uncertainties, the optimal method for determining both ToAs and their uncertainties requires further investigations.

In this study, we aim to elaborate on the findings reported in Wang et al. (2022, 2023) using 13 MSPs observed with European telescopes. The selection of these pulsars is based on their significance within the European Pulsar Timing Array (EPTA) as determined by Babak et al. (2016). Consequently, we scrutinized the achieved timing precision as a function of the chosen templates and TMMs. Additionally, we assessed the noise floor for each pulsar across the four EPTA telescopes. The remainder of this article is structured in the following manner. In Sect. 2, we provide an overview of the observations and the methodologies employed for data reduction in this study. The timing results are presented on a per pulsar basis in Sect. 3. Finally, we discuss these results and summarize our conclusions in Sect. 4.

2. Observations and analysis procedure

This paper presents observations of an array of MSPs in the L band¹ from four EPTA telescopes over 13 yr, from 2007 to 2020. The EPTA is a collaborative effort among European countries with the primary goal of directly detecting low-frequency GWs. The concept and objectives of this organization were initially outlined in detail by Stappers et al. (2006), and a more comprehensive overview of the EPTA telescopes, the observing system, and the list of primary sources can be found in the work of EPTA Collaboration (2023).

The dataset discussed in this paper was collected by four European telescopes: the Effelsberg 100 m radio telescope (EFF) in Germany, the Lovell radio telescope at the Jodrell Bank Observatory (JBO) in the UK, the Westerbork Synthesis Radio Telescope (WSRT) in the Netherlands, and the Nançay decimetric radio telescope (NCY) in France. At each telescope, pulsar observations were carried out using the latest generation of instruments, namely the PSRIX (Lazarus et al. 2016), ROACH (Bassa et al. 2016), NUPPI (Cognard et al. 2013; Liu et al. 2014), and PuMa II (Karuppusamy et al. 2008). All 13 sources analyzed in this study were observed at intervals ranging from weekly to monthly as part of the EPTA program except for NCY, which was observed more frequently. At EFF, observations were recorded across a bandwidth of 200 MHz with a center frequency of 1347.5 MHz. Data collected from both the multi-beam and single-feed 21 cm receivers were utilized for this study. Typically, each pulsar was observed once a month with an integration time of approximately 30 min. It should be noted that due to the scheduling of data acquisition, the calibration files were missing when the EFF observation data were collected. As a result, the EFF data were not polarization-calibrated. Consequently, the analysis results of the EFF data should not be interpreted as indicating any inherent disadvantages compared to data from other telescopes. Pulsars observed with the Lovell telescope at JBO were recorded using the ROACH back-end, with a center frequency of 1532 MHz and a total bandwidth of 400 MHz. NCY observations spanned approximately 10 yr, from 2011 to 2020, with an average cadence of about 5–10 days. Data were recorded at NCY with a center frequency of 1484 MHz. The total bandwidth of 512 MHz was channelized into 128 channels

¹ “ L band” refers to the frequency range of 1–2 GHz in the radio spectrum.

Table 1. Observational configurations of various EPTA telescopes pertinent to this study.

Telescope	EFF	JBO	NCY	WSRT
Pulsar back-end	PSRIX	ROACH	NUPPI	PuMa II
Frequency (MHz)	1347.5	1532	1484	1380
Bandwidth (MHz)	200	400	512	160
Average length (min)	30	20	50	30
Number of channels	128	400	128	512
Number of phase bins	1024	512	2048	512

and de-dispersed coherently. Recently, [Guillemot et al. \(2023\)](#) have deployed an improved polarization calibration scheme for the Nançay NUPPI observations. However, the data presented in our work were not calibrated with the improved scheme, and thus the profiles in our data may display variations from (time-varying) imperfect calibration parameters. WSRT observations were conducted at a center frequency of 1380 MHz across a bandwidth of 160 MHz. At WSRT, each pulsar was observed for approximately 30 min at intervals of around one month. Further details regarding the bandwidths, frequencies, and average lengths of the observations at each telescope are provided in Table 1. In Tables 2 and 3, we provide a summary of the pulsar details, the modified Julian date (MJD) ranges, and the number of observations retained after the initial data quality check.

The timing models utilized in this study are primarily based on the work of [Desvignes et al. \(2016\)](#), with updates to the reference clock and the Solar System ephemeris to “TT(BIPM2019)” and “DE438,” respectively. Since our focus is not on investigating the parameters of the timing model, we only fitted the rotational frequency and its first derivative in the analysis of templates and TMMs. However, due to variations in the dispersion measure (DM) at different subbands, we also fitted the DM, along with its first and second derivatives whenever possible, for the investigation of the system-limited noise floor (SLNF), which refers to the level of residual noise saturation observed in pulsar timing measurements when using the widest ToA bandwidths.

Prior to any further processing, it is essential to remove any radio-frequency interference (RFI) from the data. Initially, all data, except for those from JBO, underwent processing using the `clean.py` tool from the COASTGUARD package ([Lazarus et al. 2016](#)) to eliminate contaminated channel-subintegration combinations. Any remaining RFI was then manually and interactively removed through visual inspection using the PSRCHIVE ([Hotan et al. 2004b](#)) `pazi/paz` tools. For JBO data, a specific Python package called CLFD ([Morello et al. 2019](#)) was automatically applied. This package identifies and zero-weights RFI using Tukey’s rule ([Tukey 1977](#)) in a three-dimensional feature space (standard deviation, peak-to-peak difference, and absolute value of the second bin of the profile’s Fourier transform) before the data are manually inspected.

3. Results

In this section, we present the specific findings for each of the 13 pulsars in our sample and evaluate how the choice of template and TMM affects data from different EPTA telescopes. In addition, we apply the SLNF analysis and determine the optimal ToA bandwidth. We begin by discussing individual pulsars with distinctive results. For pulsars that exhibit similar outcomes, we provide a collective description at the end of this section.

Given that we have previously highlighted the subpar performance of templates from the single brightest observation and of the GIS TMM (see [Wang et al. 2022](#)), and because we consistently observe the same pattern from the majority of our current sample, we refrain from reiterating the comparison and description of the single template and GIS TMM analysis in the subsequent text, unless a specific and unique result emerges for a particular pulsar, telescope, and back-end combination.

3.1. PSR J0030+0451 and PSR J2010–1323

PSR J0030+0451 and PSR J2010–1323 are two of the four isolated MSPs in our sample. PSR J0030+0451 is one of the closest known MSPs in the Galaxy, situated at a distance of ~ 325 parsecs ([Lommen et al. 2000](#); [Arzoumanian et al. 2018](#)). Due to the likelihood of strong influence from solar wind at low solar elongations, observations taken when the angular separation from the Sun is less than 5 degrees (see [Tiburzi et al. 2021](#); [Verbiest et al. 2016](#), for detailed discussions) have been excluded from our analysis. PSR J2010–1323 was discovered by [Jacoby et al. \(2007\)](#) during the high Galactic-latitude survey.

At some of the EPTA telescopes, only a small number of observations were made for these two pulsars, and thus our results must be interpreted with caution for those datasets. PSR J0030+0451 was observed only 35 times at WSRT, while PSR J2010–1323 had 34 observations with EFF. In addition, with mean S/N values of 25.30 and 50.02, respectively, both of these pulsars have low S/N. The timing analysis and SLNF fitting are significantly impacted by the small sample size and low S/Ns. The outlier rejection strategy, as defined in [Wang et al. \(2022\)](#), which primarily relies on the S/N criterion, exacerbates these issues, resulting in an even smaller sample size.

In the cases of EFF and WSRT, the results of the template and TMM choices lead to two completely opposing outcomes. For PSR J0030+0451 at WSRT, while most other combinations yield consistent and intuitive results, the added template combined with PGS results in an almost halved RMS of the residuals compared to FDM. The superiority of the added template in this case is likely due to imperfections in both the smoothing algorithms and von Mises functions fitting process, which remove certain features (as can be seen in Fig. A.1) in the added template. In other words, the smoothed and analytic templates do not always accurately represent the true profile. However, the situation for PSR J2010–1323 at EFF is completely reversed, with the added template combined with PGS resulting in an almost doubled RMS compared to FDM. Both combinations are outperformed by other options. After inspecting the profile differences, no distinct features were found between the added template and the analytic template, indicating that the noise-free template can accurately represent the profile. Therefore, we attribute the improvement to the removal of random noise.

Regarding the SLNF analysis, the fitting for these two special cases fails due to the limited number of ToAs and low S/N. In cases with small ToA bandwidths, a larger proportion of ToAs are eliminated by our outlier rejection scheme, leading to biased RMS values. Aside from the specific cases mentioned above, when these two pulsars are observed with other telescopes, the results align with the conclusions drawn in [Wang et al. \(2022\)](#): added and smoothed templates are typically preferred over some more commonly applied alternatives, and the FDM method is generally superior to or competitive with other algorithms. At NCY, PSR J0030+0451 is the most frequently observed pulsar in our dataset, with a cadence of 5 days. The RMS values of PSR J0030+0451 obtained with NCY decrease more rapidly than

Table 2. Overview of the pulsar observations.

PSR	EFF		JBO		NCY		WSRT		Data span (yr)
	MJD range	N_{obs}	MJD range	N_{obs}	MJD range	N_{obs}	MJD range	N_{obs}	
J0030+0451	55633–58859	55	56095–58495	101	55803–59152	706	56080–57113	35	9.6
J0613–0200	55600–58620	91	56253–58494	83	55817–58851	318	54337–57096	101	12.4
J1012+5307	55633–59083	103	55655–58495	144	55816–59288	799	54165–57153	83	14.0
J1022+1001	55660–58769	109	55600–58131	184	55839–58853	280	54154–57153	80	12.9
J1024–0719	55633–59223	79	55674–59120	158	55891–59153	393	54287–57118	75	13.3
J1600–3053	55661–58621	60	56255–58559	84	55809–58852	624	54155–57131	62	12.9
J1640+2224	55633–58620	94	55653–58467	95	55826–58850	369	54155–57101	86	12.9
J1744–1134	55628–59161	87	55670–58560	94	55805–59212	213	55792–57154	52	9.4
J1857+0943	55633–58824	64	56011–58559	121	55800–58837	235	54337–57177	81	12.4
J1909–3744	–	–	–	–	55812–59277	515	–	–	9.5
J1918–0642	55633–59223	76	55652–58838	86	55874–59138	169	54490–57190	86	13.0
J2010–1323	56206–59223	34	55657–58494	96	55850–59149	298	54492–57148	75	13.0
J2317+1439	55633–58824	58	55662–58495	83	55850–58849	701	54490–57117	78	11.9

Notes. Shown are the range of modified Julian dates (MJDs) for pulsars observed by different telescopes, with the number of obtained observations (N_{obs}), and the overall data span for each pulsar-telescope combination.

Table 3. Overview of the pulsars' properties.

PSR	P (ms)	\dot{P} (10^{-20})	P_b (days)	DM (cm^{-3} pc)	$S_{1.4}$ (mJy)	Ref.
J0030+0451	4.865	1.02	–	4.33	0.6	Lommen et al. (2000); Verbiest et al. (2016)
J0613–0200	3.062	0.96	1.20	38.79	1.7	Lorimer et al. (1995); Verbiest et al. (2016)
J1012+5307	5.26	1.71	0.60	9.02	2.8	Nicastro et al. (1995); Lange et al. (2001)
J1022+1001	16.453	4.33	7.81	10.25	6.1	Camilo et al. (1996); Desvignes et al. (2016)
J1024–0719	5.162	1.86	–	6.49	1.5	Bailes et al. (1997); Manchester et al. (2013)
J1600–3053	3.598	0.95	14.35	52.33	2.5	Jacoby et al. (2007); Ord et al. (2006); Manchester et al. (2013)
J1640+2224	3.163	0.28	175.46	18.43	0.4	Foster et al. (1995); Desvignes et al. (2016)
J1744–1134	4.075	0.86	–	3.14	1.0	Bailes et al. (1997)
J1857+0943	5.362	1.78	12.33	13.30	3.3	Segelstein et al. (1986); Verbiest et al. (2016)
J1909–3744	2.95	1.40	1.53	10.39	1.0	Jacoby et al. (2003); Desvignes et al. (2016)
J1918–0642	7.646	2.57	10.91	26.55	0.6	Edwards & Bailes (2001b); Janssen et al. (2010)
J2010–1323	5.223	0.48	–	22.16	1.6	Jacoby et al. (2007)
J2317+1439	3.445	0.24	2.46	21.90	4.0	Camilo et al. (1993); Verbiest et al. (2016)

Notes. The pulsar's rotation period (P), spin-down (\dot{P}), orbital period (P_b), DM, and flux density at 1.4 GHz ($S_{1.4}$) are given.

with other telescopes as a function of ToA bandwidth, and the asymptotic RMS limit, as low as 1.18 μs , appears to be reached with a ToA bandwidth of about 1 GHz. With an SLNF of 1.45 μs , which is achievable with the existing backend, NCY is likewise the best-performing telescope for PSR J2010–1323. Extending the available bandwidth is not expected to significantly increase timing precision.

3.2. PSR J1012+5307

PSR J1012+5307, discovered by Nicastro et al. (1995), resides in a 14.5-h binary system with an extremely low-mass helium-core white dwarf companion, estimated to be around 0.16 M_{\odot} (van Kerkwijk et al. 1996; Callanan et al. 1998). Leveraging data from EFF and JBO, Lange et al. (2001) established a precise timing model for PSR J1012+5307, and determined the detailed three-dimensional velocity motion for this binary system through a combination of radio timing and optical observations. A subsequent study by Lazaridis et al. (2009) revisited PSR J1012+5307, further refining the three-dimensional velocity

information with 15 yr of EPTA data. One noteworthy outcome was the establishment of a new upper limit for intrinsic eccentricity, which is less than 8.4×10^{-7} , positioning it as one of the smallest values recorded for a binary system.

PSR J1012+5307 is among the most frequently observed, as well as having one of the longest data-spans pulsars within the EPTA, particularly with the NCY telescope. While this pulsar is observed approximately once per month at EFF, JBO, and WSRT, NCY observes this pulsar roughly every 6 days. In the case of EFF data, this pulsar exhibits a significantly high reduced chi-square χ_r^2 , likely attributable to profile variations stemming from uncalibrated polarization (refer to Fig. 1). Except for this telescope, the conclusions drawn are consistent. The combination of the added template and FDM method produces more reliable and precise ToAs. PGS may yield a slightly improved RMS value, but it substantially elevates the χ_r^2 . The SLNF fitting results depicted in Fig. C.1 indicate that this pulsar's performance observed with EFF is suboptimal, and increasing the ToA bandwidth does not significantly reduce the RMS. This is most likely due to our failure to calibrate the EFF data.

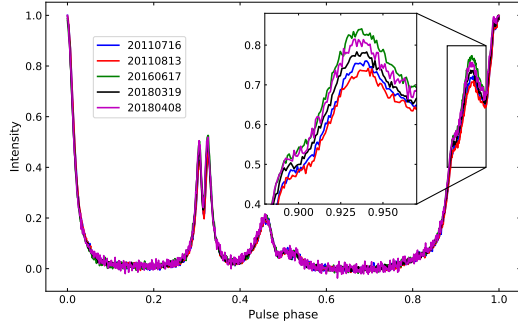


Fig. 1. Example of profile variations of PSR J1012+5307 at EFF, with each color representing a peak-normalized observation taken at a certain date, as indicated in the legend. A zoomed-in view of the leading component of the main pulse is shown in the subplot. These variations are not found in other telescopes, and thus we attribute them to the fact that we did not polarization-calibrate the EFF data.

3.3. PSR J1022+1001

PSR J1022+1001 is an MSP with a spin period of 16.45 ms, discovered concurrently by the Princeton-Arecibo Declination-Strip Survey (Camilo et al. 1996) and the Green Bank Northern Sky Survey (Sayer et al. 1997). It is in a binary system with a 7.8-day orbit around a CO white dwarf companion. The average pulse profile of PSR J1022+1001 in the L band displays two prominent peaks. While the trailing peak component exhibits high linear polarization, the leading component does not show strong polarization (Hotan et al. 2004a). The pulse profile of PSR J1022+1001 is observed to undergo significant intrinsic variations with frequency (e.g., see Fig. 2 and Ramachandran & Kramer 2003; Liu et al. 2015). Due to strong scintillation in the L band, the intensity ratio of its leading and trailing peaks varies noticeably over time. While Kramer et al. (1999) argue that the profile variations in the integrated pulse profile cannot be attributed to instrumental and propagation effects, van Straten (2013) find that instrumental polarization artifacts play a major role in the profile instability. van Straten (2013) propose a new polarimetric calibration method, referred to as measurement equation template matching (METM), to address this instability. Based on this novel method, timing results from 7.2 yr of data of this pulsar from the Parkes radio telescope show that the standard deviation of arrival time residuals derived from METM-calibrated data is two times smaller than with conventional calibration methods.

In addition to profile variations, this pulsar is also located at a very low ecliptic latitude (as low as -0.06 deg), suggesting that the pulsar signals could be significantly influenced by solar wind. Generally, dispersion in the solar wind leads to DM fluctuations and an incomplete model of those variations can introduce delays of up to a few microseconds in the ToAs (Tiburzi et al. 2021). As only a small fraction of the observations were taken within five degrees of the Sun, we excluded the observations taken when the Solar elongation was less than five degrees.

Within our dataset, we observe significant profile variations across the observed frequency band, especially in EFF, JBO, and NCY data. In WSRT data, the profile variations are still noticeable but not as pronounced as for other telescopes (presumably due to the narrower bandwidth). Due to these profile variations, ToAs obtained with these telescopes exhibit high χ_r^2 values. However, despite the issue of high χ_r^2 values, the timing precision at all telescopes remains consistent with each other.

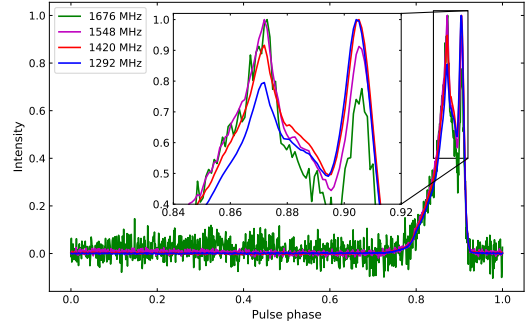


Fig. 2. Example of profile variation of PSR J1022+1001 obtained with NCY data as a function of frequency. The entire band is divided into four subbands with a subband width of 128 MHz. The data underwent polarization calibration, and each subband profile was peak-normalized.

For this pulsar, the WSRT data exhibits the lowest SLNF at approximately $1.22 \mu\text{s}$. Additionally, the RMS decreases gradually as a function of ToA bandwidth, similar to JBO and NCY. However, the EFF data behave differently. The increase in ToA bandwidth does not lead to as significant an improvement in RMS as observed in the other datasets, which might be attributed to the lack of polarization calibration for those data.

3.4. PSR J1600–3053

PSR J1600–3053 is relatively bright for an MSP, especially considering its large DM value as shown in Table 3. Discovered in the Swinburne high Galactic-latitude pulsar survey by Jacoby et al. (2007), this binary MSP has since been well timed with very high precision, owing to its inherently narrow pulse profile and high brightness.

During our analysis, we observed that this pulsar is affected by strong DM variations from the end of 2014 onward. Figure 3 displays the DM time series for PSR J1600–3053, determined with the NCY dataset. We calculated the DM offset using the method described in Donner et al. (2020) and introduced the “-dmo” flag in the ToA files to correct for these variations. After correcting for the effects of DM variations, the results for templates and TMMs are shown in Fig. B.1. We see that for this pulsar the added template works very effectively. The most significant improvement is seen in the WSRT data, where the combination of the added template with the PGS method leads to a substantial decrease in RMS and brings the χ_r^2 value in line with the data from other telescopes.

For this pulsar, the profile difference (see Fig. A.1), between the added and smoothed, analytic templates, is noticeable, particularly with the NCY data. These characteristics were not alleviated by employing other smoothing algorithms or fitting with additional Gaussian functions, which could explain the modest deterioration in the performance of the smoothed and analytic templates.

Regarding the SLNF, it is evident that the RMS decreases noticeably as the ToA bandwidth increases. JBO and NCY exhibit a similar rate of decrease, although their final SLNF values differ significantly, being approximately 566 ns and 900 ns, respectively. Another noteworthy observation is that the SLNF fit for WSRT data failed with this pulsar. We attribute this to the low S/N, with a mean S/N of around 20.46 for the observations. Low S/N levels in observations could lead to underestimated ToA uncertainties and result in extra outliers, particularly in frequency-resolved timing where the S/N is diminished with a

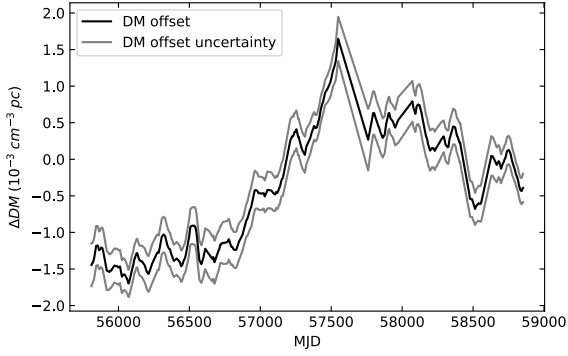


Fig. 3. DM time series for PSR J1600–3053 obtained with NCY data, with a reference DM of $52.321 \text{ cm}^{-3} \text{ pc}$.

decrease in ToA bandwidth. Moreover, low S/N levels can lead to a significant number of ToAs being removed by the outlier rejection scheme. For a ToA bandwidth of 20 MHz or less, over 40% of measurements can be eliminated for this pulsar.

3.5. PSR J1909–3744

PSR J1909–3744 is the fastest-spinning MSP in our dataset, with a spin period of 2.95 ms, orbiting in a 1.53-day binary system. It is prominent for its extremely stable pulse profile characterized by a narrow peak and a small pulse duty cycle of approximately 1.5%. Due to its low declination, in the EPTA it is exclusively observed by the NCY telescope. Using data from NCY, Liu et al. (2020) managed to achieve a timing precision of approximately 100 ns on a 15-yr timescale. Notably, studies by Shannon et al. (2014) and Lentati & Shannon (2015) have revealed that this pulsar exhibits an unexpectedly low level of jitter noise (around 10 ns), indicating significant potential for further improvements in timing precision.

In light of the presence of achromatic red noise and DM noise components (Liu et al. 2020; Chalumeau et al. 2022), the initial RMS and χ_r^2 values are relatively high, standing at 0.241 ms and approximately 25, respectively. To address this, we applied the red noise model established in Liu et al. (2020), resulting in substantial improvements in the timing solution. The RMS and χ_r^2 dropped significantly to approximately 70 ns and 1.8, respectively.

Figure 4 illustrates that for this exceptionally precisely timed pulsar, we observed no notable differences after exploring various combinations of templates and TMMs, when we excluded the obviously inaccurate behavior of GIS. Additionally, we find that the discrepancy between the mean error bars was negligible. In terms of the ToA bandwidth analysis, the noise fitting curve, as depicted in Fig. 5, aligns with the RMS trend as a function of ToA bandwidth remarkably well. The fitting also suggests that solely increasing the ToA bandwidth of the NCY backend would not result in a significant enhancement in timing precision. We estimated the SLNF for this pulsar to be approximately 54 ns.

3.6. PSR J1918–0642

PSR J1918–0642 is an MSP with a spin period of 7.65 ms, located in a binary system with an orbital period of 10.91 days. It was initially discovered by Edwards & Bailes (2001a) in a survey of intermediate Galactic latitudes. Janssen et al. (2010) conducted a detailed study of this pulsar using 7.5 yr of EPTA observations and updated its timing model. They observed significant flux density modulation in PSR J1918–0642 due to refractive and

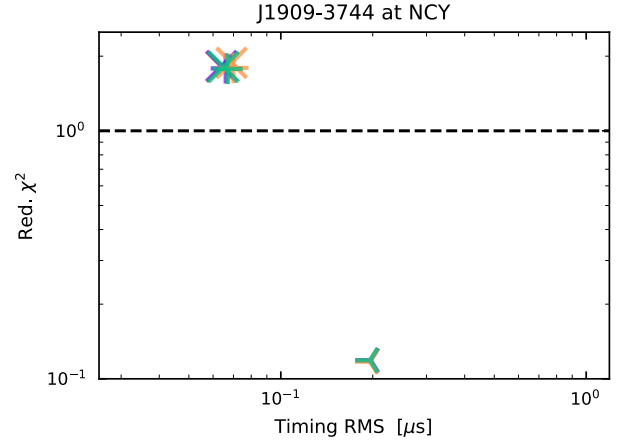


Fig. 4. χ_r^2 and residual RMS for PSR J1909–3744 obtained with NCY data as a function of the selected TMM and template. In the plot, the “+,” “x,” and “tri-left” symbols represent the three TMMs, FDM, PGS, and GIS, respectively. The different colors represent the different templates: added (blue), analytic (cyan), single (orange), and smoothed (olive). As seen in this plot, the GIS technique, compared with other techniques, produces much larger RMS values and inaccurate ToA error estimates.

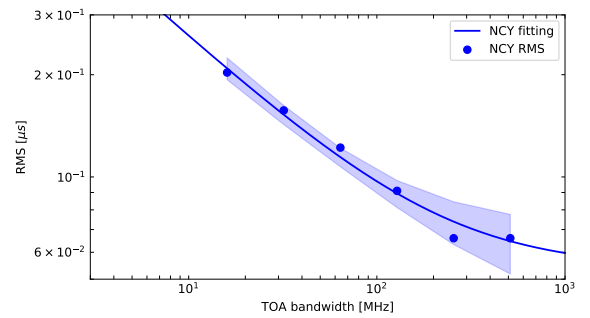


Fig. 5. SLNF estimates for PSR J1909–3744 when the added template with the FDM TMM is used to generate the ToA with the available bandwidth obtained with NCY data divided into the respective number of channels. The shaded area along the fitted curve is the interquartile range for the fit, representing the error bound on the estimated SLNF.

diffractive scintillation. Subsequent studies by Desvignes et al. (2016) and Fonseca et al. (2016) using 12.5 yr of EPTA data and 9 yr of NANOGrav data, respectively, successfully detected the Shapiro delay in this system.

In our analysis of PSR J1918–0642, we find that the results obtained from JBO data align with the conclusions of Wang et al. (2022). However, for EFF, NCY, and WSRT ToAs, we observed that using added templates with the FDM method resulted in approximately 10% higher RMS compared to using PGS with added templates. This difference was most pronounced at WSRT, where PGS led to a 15% improvement in both RMS and χ_r^2 .

3.7. Other pulsars

The results for the remaining six pulsars in our dataset are consistent with the conclusions drawn in Wang et al. (2022). The FDM method consistently outperforms or is on par with other TMMs. Additionally, the added, smoothed, and analytic templates yield comparable results. In certain cases, such as PSR J0613–0200, using the PGS method with added templates results in a slightly lower RMS, but at the expense of increased

χ_r^2 . This suggests that for PSR J0613–0200, PGS tends to underestimate the ToA uncertainty. Furthermore, PSR J1640+2224 and PSR J2317+1439 demonstrate that using the added template in combination with PGS leads to suboptimal determination of ToAs and their uncertainties, as evidenced by noticeably higher RMS and χ_r^2 values.

For PSR J1024–0719, PSR J1744–1134, and PSR J1857–0642, the results are consistent across most combinations of templates and TMMs yield similar outcomes. We also observe that ToAs determined by FDM and PGS are the most reliable in terms of ToA uncertainties. FDM tends to perform significantly better than PGS, particularly when templates are not noise-free, resulting in lower χ_r^2 values.

Regarding the SLNF, different pulsars and telescopes yield markedly different results. With a few exceptions, such as PSR J1744–1134, where EFF and WSRT achieve the best SLNF at approximately 300 ns, NCY consistently attains the lowest SLNF. Additionally, EFF exhibits the sharpest decreasing trend, suggesting that adopting a suitable ToA bandwidth can help reduce system noise and bring it closer to the SLNF limit.

4. Conclusions

This study focused on evaluating the impact of template and TMM choices on pulsar timing precision for a large sample of EPTA pulsars. Our findings demonstrate that the GIS method is not suitable for achieving high-precision timing. Regardless of the template used, GIS consistently resulted in relatively high RMS values for the timing residuals. Additionally, GIS tended to overestimate the ToA uncertainties, leading to seemingly over-fitted timing residuals. Moreover, FDM generally outperforms the PGS algorithm with a few exceptions, typically due to a limited number of ToAs and low-S/N of observations (e.g., PSR J0030+0451 at WSRT and PSR J2010–0719 at EFF) or underestimated ToA uncertainties (e.g., PSR J1640+2224 and PSR J2317+1439). Based on our results, we recommend using FDM as the default method for high-precision pulsar timing analysis.

In terms of templates, we highly recommend using data-derived and smoothed templates. While these templates may be susceptible to the “self-standardizing” issue discussed in Hotan et al. (2005), we did not observe any evidence of this in our templates. It is, in this regard, worth noting that the self-standardizing problem tends to diminish when the S/N of individual observations is above 25 (Hotan et al. 2005). Given that most of our observations have a high S/N, this issue has little impact on our investigation. Furthermore, with ongoing hardware upgrades and the development of RFI mitigation algorithms, self-standardizing is becoming less of a concern for PTA analysis. Therefore, we recommend using added templates as the default for current and future PTA ToA determination.

When categorizing templates and TMMs by telescope, several conclusions become evident. Firstly, all templates and TMMs yield stable results with NCY data, with negligible differences in RMS or χ_r^2 values. This suggests that the choice of template or TMM has a minimal impact on NCY data. JBO data show a similar pattern to NCY data, except that PGS combined with added templates can sometimes lead to unsatisfactory results, particularly for PSR J0030+0451, PSR J1640+2224, and PSR J2317+1439, which tend to have significantly high χ_r^2 values. With WSRT data, in contrast to the added template, analytic and smoothed templates yield consistent timing results. While PGS combined with added templates may result in high χ_r^2 values for most pulsars, PSR J0030+0451 and PSR J1600–3053

show significant RMS improvements with the added template. For EFF, the added template combined with FDM is generally preferred, as it tends to yield either better RMS and χ_r^2 values or consistent results. However, there are two exceptions, PSR J1744–1134 and PSR J2010–1323, for which smoothed and analytic templates provide consistent timing results, both superior to the added template.

Overall, our findings suggest that with the improvement in telescope sensitivity, template and TMM choices will have a smaller impact on pulsars currently timed for PTA purposes. However, for the time being, it is important to carefully consider TMM and template selection, especially for telescopes with lower sensitivity. Based on available data and experimental comparisons, we strongly recommend using a combination of data-derived and smoothed templates, along with the FDM method, for determining ToAs in the current EPTA project. It is important to note that special cases should not be overlooked, and different combinations should be compared in practical work.

By analyzing the timing RMS variations as a function of the ToA bandwidth and comparing our calculated SLNF with the jitter noise reported in other studies (e.g., Parthasarathy et al. 2021; Lam et al. 2019), we can estimate the current SLNF for various telescope systems and pulsars. It is worth mentioning that the EFF results in this paper are all worst-case estimates and that proper calibration is expected to significantly improve these results, although this improvement may be significantly pulsar-dependent. Given that the jitter levels observed in previous research are notably lower than the SLNF, we can conclude that pulse phase jitter noise is not a limiting factor for current EPTA timing accuracy.

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Appendix A: Pulse profiles and profile differences

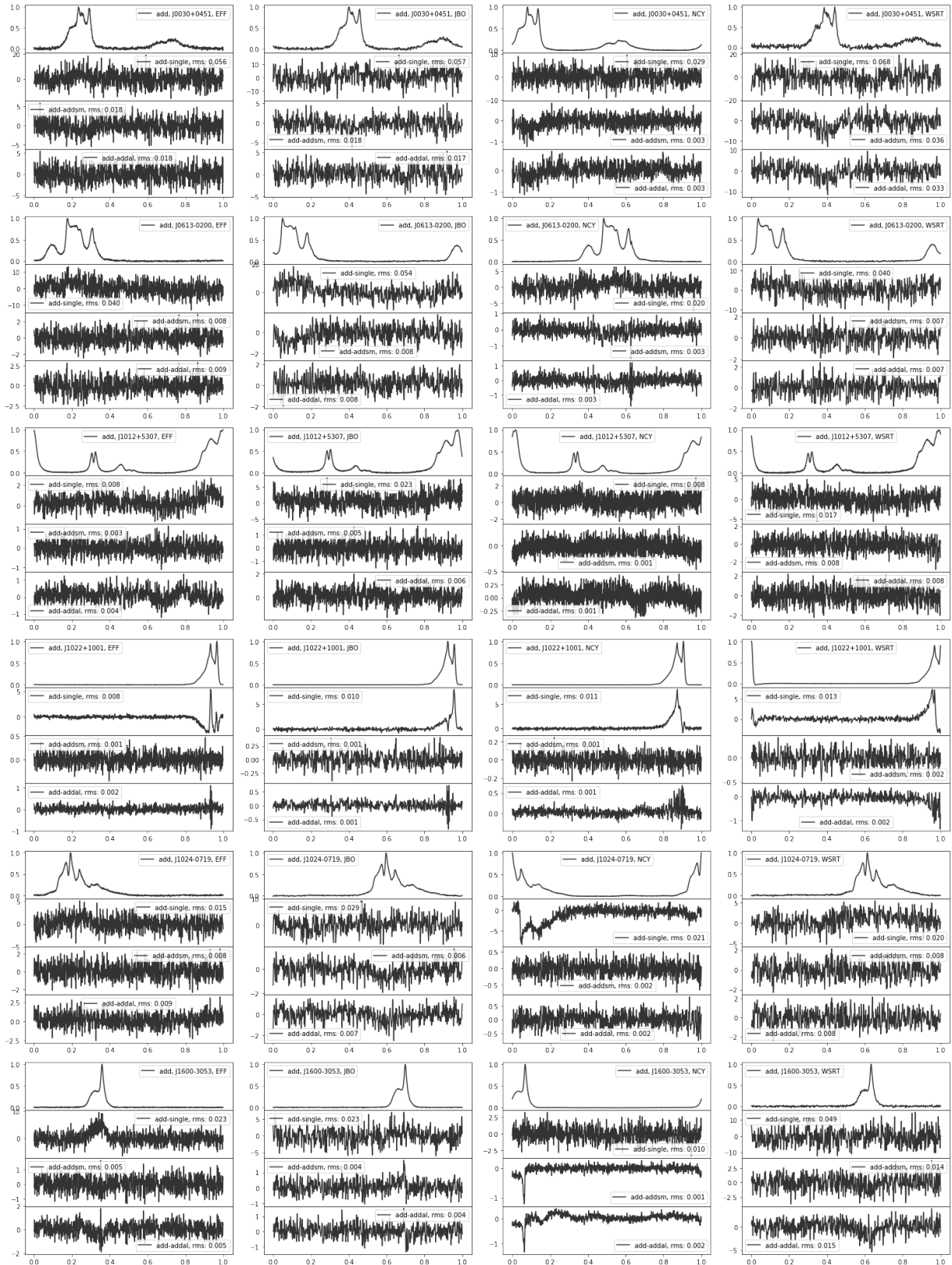


Fig. A.1: Pulse profiles of the added templates and profile differences between added and single, analytic, and smoothed templates. From top to bottom, PSR J0030+0451, PSR J0613-0200, PSR J1012+5307, PSR J1022+1001, PSR J1024-0719, and PSR J1600-3053. Within each subplot, the profile of each pulsar and the difference between added and single, smoothed, and analytic templates are shown from top to bottom. In each row, the EFF, JBO, NCY, and WSRT data are shown from left to right.

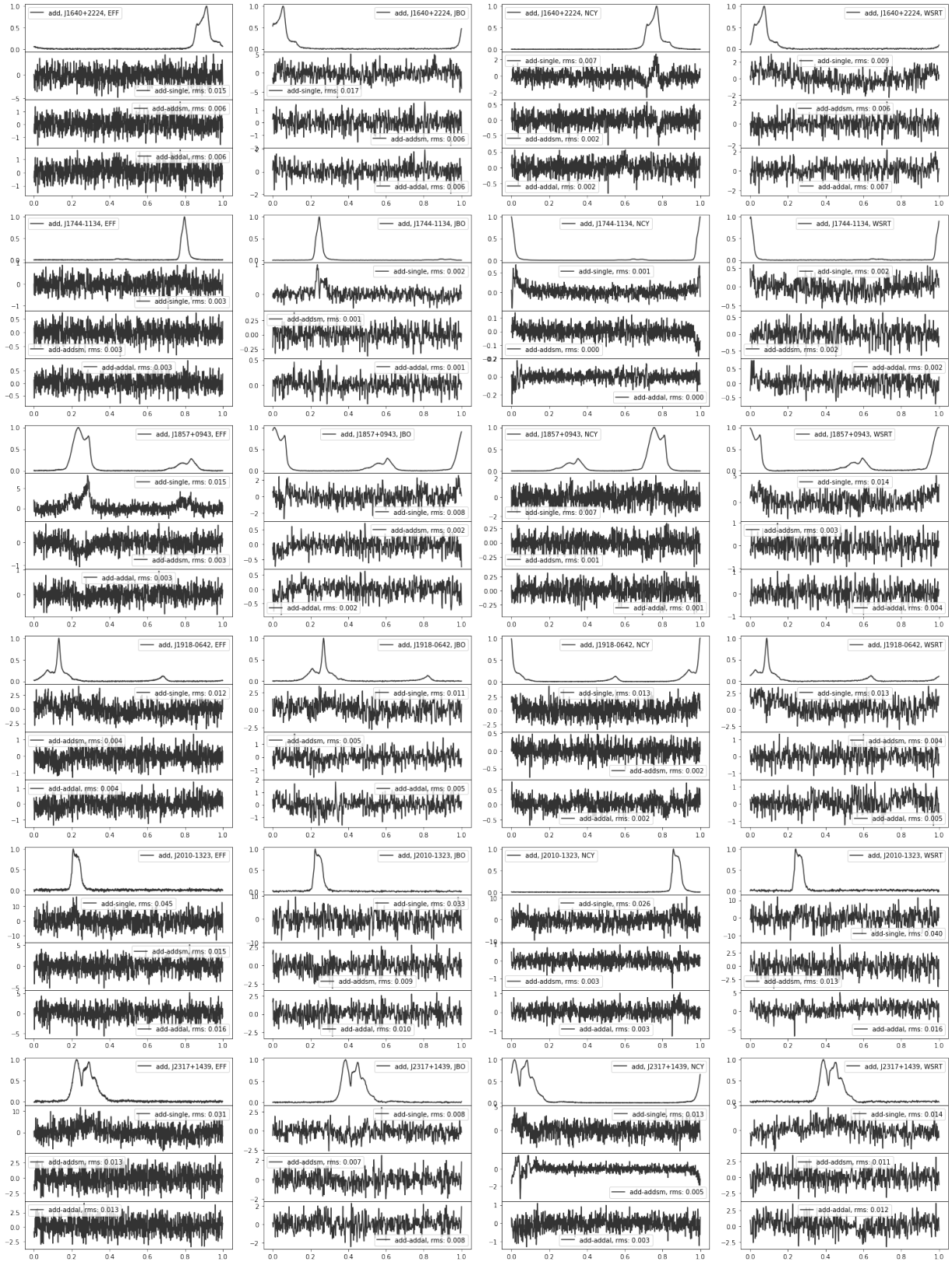


Fig. A.1: continued: (PSR J1640+2224, PSR J1744–1134, PSR J1857+0943, PSR J1918–0642, PSR J2010–1323, and PSR J2317+1439).

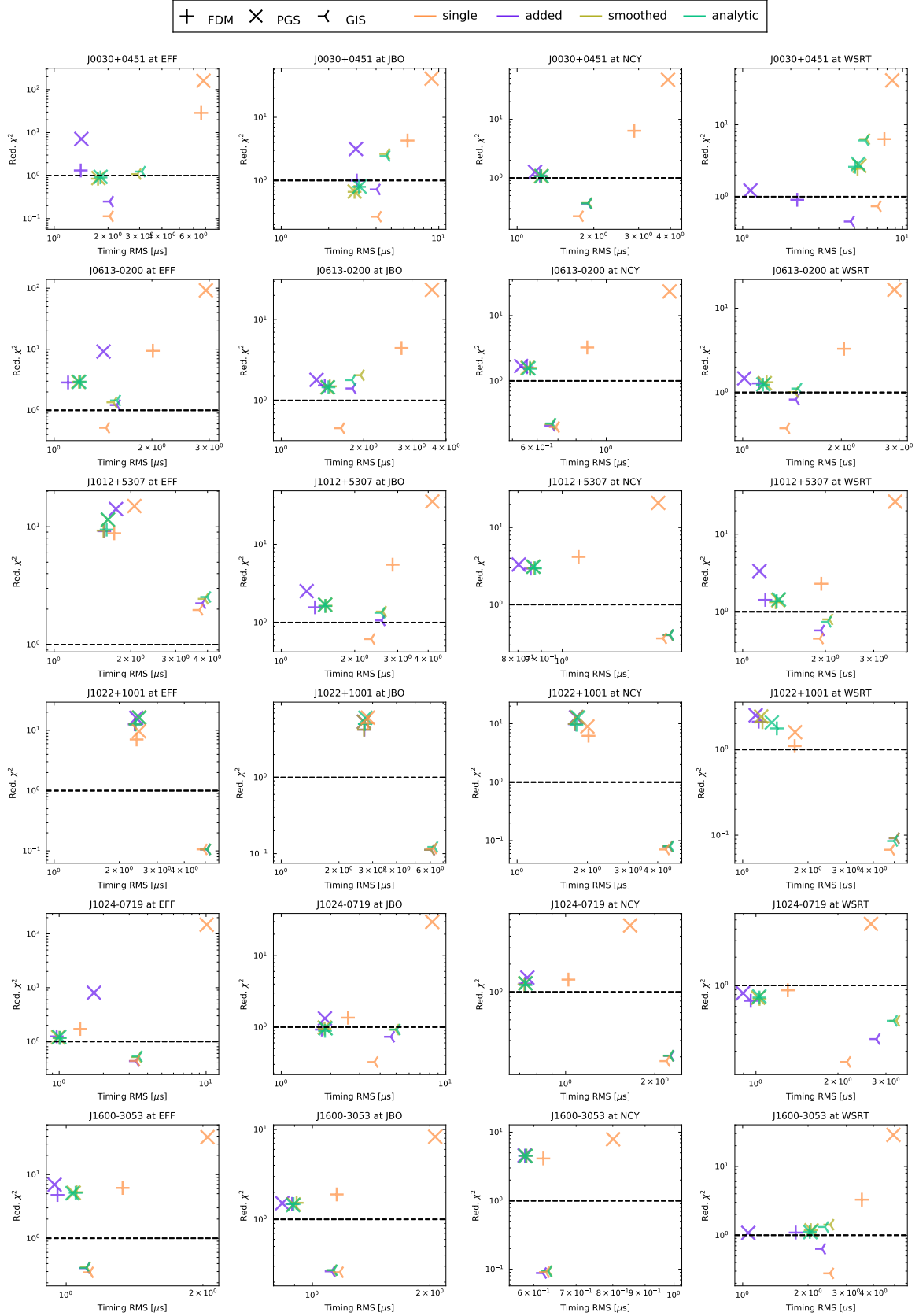
Appendix B: χ_r^2 and residual RMS

Fig. B.1: χ_r^2 and residual RMS for the same pulsars as in Fig A.1, as a function of the selected TMM (marker shape) and template (marker color) and the different telescopes.

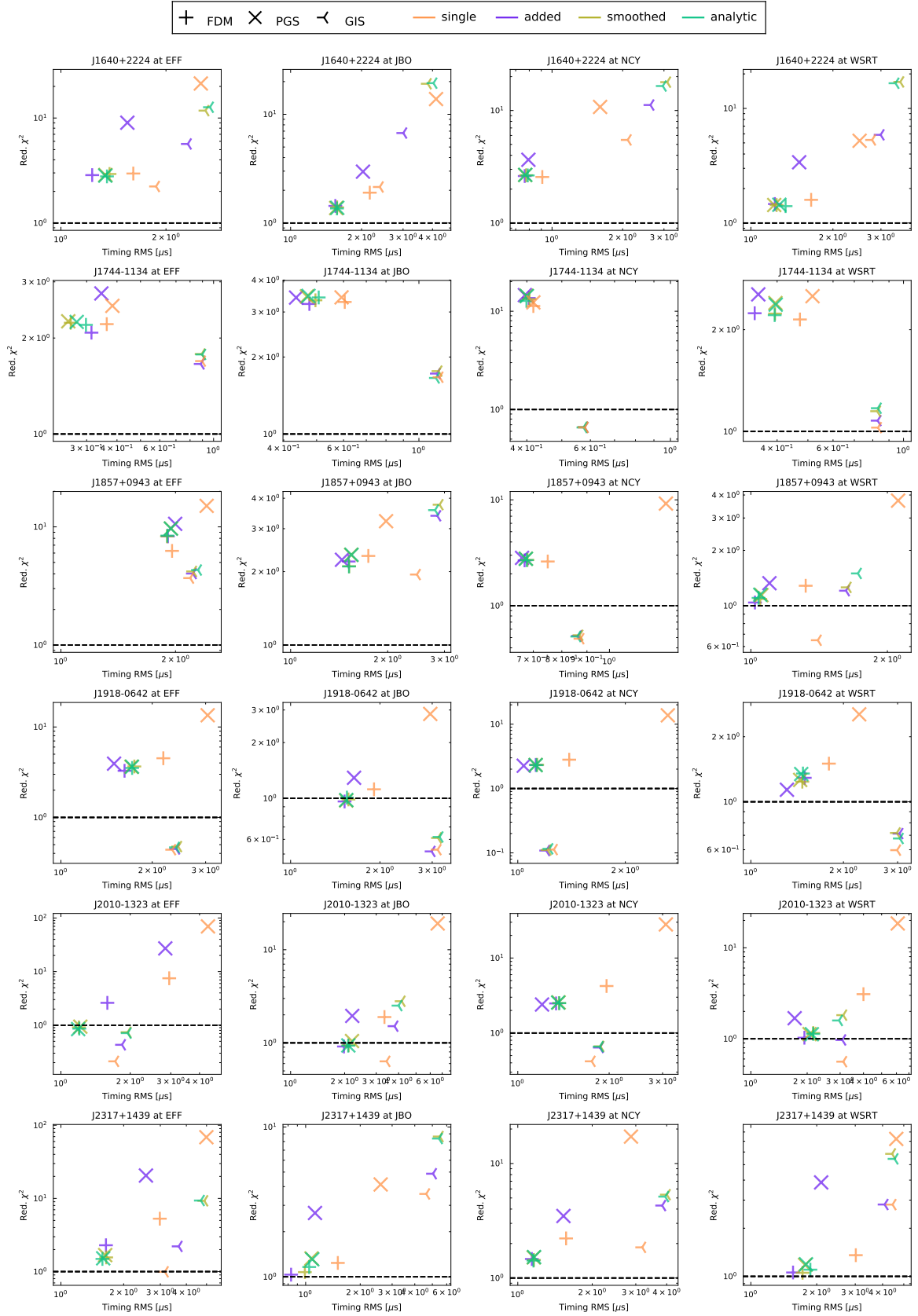


Fig. B.1: continued: (PSR J1640+2224, PSR J1744–1134, PSR J1857+0943, PSR J1918–0642, PSR J2010–1323, and PSR J2317+1439).

Appendix C: SLNF estimates for the remaining 12 pulsars

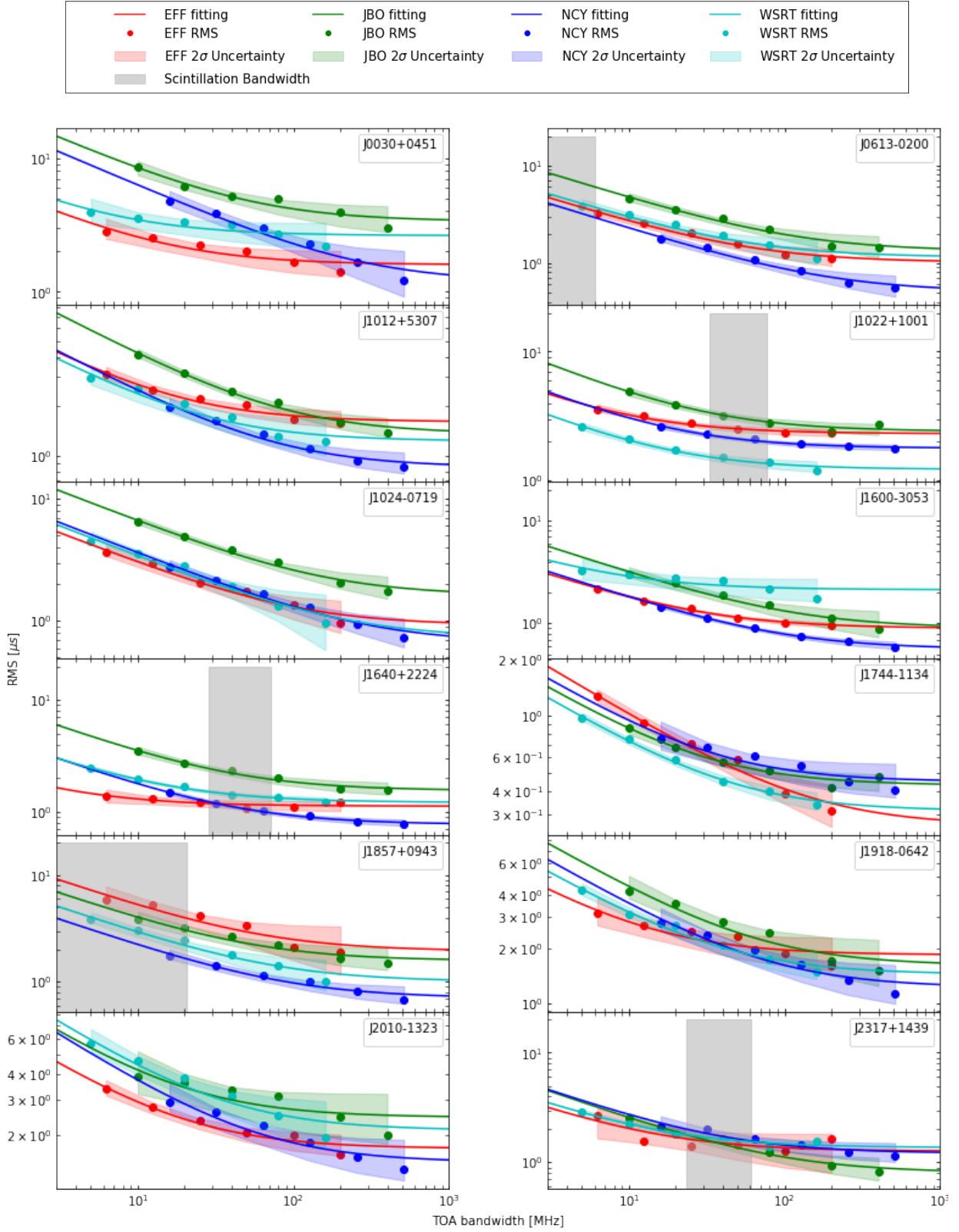


Fig. C.1: SLNF estimates for the remaining 12 pulsars.